

## The CHAOS-4 geomagnetic field model

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### SUMMARY

We present CHAOS-4, a new version in the CHAOS model series, which aims to describe the Earth's magnetic field with high spatial and temporal resolution. Terms up to spherical degree of at least  $n = 85$  for the lithospheric field, and up to  $n = 16$  for the time-varying core field are robustly determined.

More than 14 yr of data from the satellites Ørsted, CHAMP and SAC-C, augmented with magnetic observatory monthly mean values have been used for this model. Maximum spherical harmonic degree of the static (lithospheric) field is  $n = 100$ . The core field is expressed by spherical harmonic expansion coefficients up to  $n = 20$ ; its time-evolution is described by order six splines, with 6-month knot spacing, spanning the time interval 1997.0–2013.5. The third time derivative of the squared radial magnetic field component is regularized at the core–mantle boundary. No spatial regularization is applied to the core field, but the high-degree lithospheric field is regularized for  $n > 85$ .

CHAOS-4 model is derived by merging two submodels: its low-degree part has been derived using similar model parametrization and data sets as used for previous CHAOS models (but of course including more recent data), while its high-degree lithospheric field part is solely determined from low-altitude CHAMP satellite observations taken during the last 2 yr (2008 September–2010 September) of the mission. We obtain a good agreement with other recent lithospheric field models like MF7 for degrees up to  $n = 85$ , confirming that lithospheric field structures down to a horizontal wavelength of 500 km are currently robustly determined.

**Key words:** Inverse theory; Magnetic anomalies: modelling and interpretation; Rapid time variations; Satellite magnetics.

### 1 INTRODUCTION

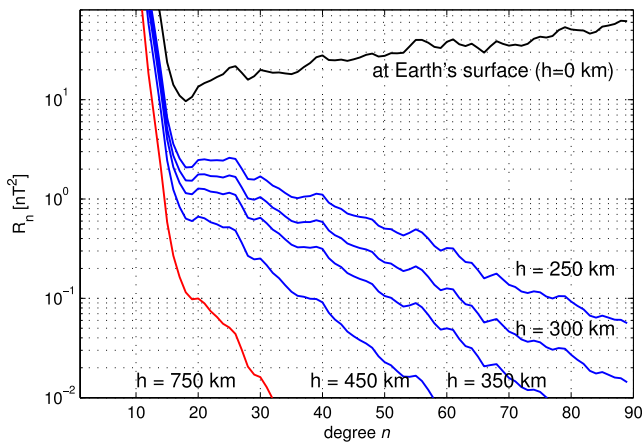
More than 14 yr of continuous magnetic field measurements collected by the Ørsted and CHAMP satellites present outstanding opportunities to investigate the Earth's magnetic field, both regarding rapid changes of the core field (secular variation) and the determination of the magnetic field due to lithospheric magnetization.

Of special interest for deriving the small-scale structure of the lithospheric field are the measurements obtained during the final months of the CHAMP satellite mission, when the satellite altitude was below 300 km. The benefit of low-altitude CHAMP observations for determining the lithospheric field is obvious from Fig. 1, which shows the spatial power spectrum of the geomagnetic field at various altitudes. The black curve is the lithospheric power at Earth's surface (according to the CHAOS-4 model presented in this arti-

cle), while the coloured curves present spectra at various altitudes of CHAMP (blue curves) and of the Ørsted satellite (red curve).

Let us assume that measuring conditions (including instrument errors, external field disturbances and spatio-temporal sampling) are such that magnetic field structures of squared amplitudes larger than  $0.1 \text{ nT}^2$  can be resolved. Under this assumption only features up to spherical harmonic degree  $n = 20$  can be determined from data obtained by a satellite at an altitude of 750 km (for example, Ørsted). In contrast, data from CHAMP at the beginning of its mission, when altitude was about 450 km, would allow for a determination of the lithospheric field up to  $n = 40$ . Data taken at 300 km altitude (which was the mean altitude of CHAMP at the beginning of 2010) would allow the determination of models up to  $n = 60$ . Finally, taking advantage of the very-low altitude data from the last weeks of CHAMP mission lifetime, as is done in this paper, would allow the lithospheric field up to  $n = 80$  to be resolved. Note that a threshold value of  $0.1 \text{ nT}^2$  has been chosen here for illustration purposes only; we do not claim that this is the actual accuracy of the magnetic satellite data.

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**Figure 1.** Spatial power spectrum of the geomagnetic field at Earth's surface (black curve) and at various altitudes of CHAMP (blue). Also shown is the spectrum at a mean altitude of the Ørsted satellite of 750 km (red).

As described in the review article by Thébault *et al.* (2010), two complementary philosophies are currently adopted when deriving spherical harmonic models of the Earth's lithospheric field. In the sequential approach, *a-priori* models of all known magnetic field contributions, except for the lithospheric field, are subtracted from the data, followed by a careful data selection and application of empirical corrections. The different versions in the MF model series derived by Stefan Maus and co-workers (*cf.* Maus *et al.* 2002, 2006, 2007, 2008) are examples of models determined using this approach. They are derived from CHAMP observations which have been along-track filtered after removal of *a priori* models of the core field and of the ocean tidal magnetic signal. MF6, the most recent published model version (Maus *et al.* 2008), also includes line levelling between adjacent satellite tracks, which minimizes the variance between close encounters. This model formally describes the lithospheric field up to spherical harmonic degree  $n = 120$  (corresponding to 333 km horizontal wavelength), but coefficients above  $n > 80$  are damped (regularized) and thus probably not robustly resolved. A new (unpublished) version, MF7 (Maus 2010), describes the field formally up to  $n = 133$ .

Contrary to this serial approach, the comprehensive approach aims to solve simultaneously for all major internal and external field contributions. Examples of models derived using this approach are the CM models (e.g. Sabaka *et al.* 2004), and, using more recent satellite data, the GRIMM models (Lesur *et al.* 2008; Lesur *et al.*

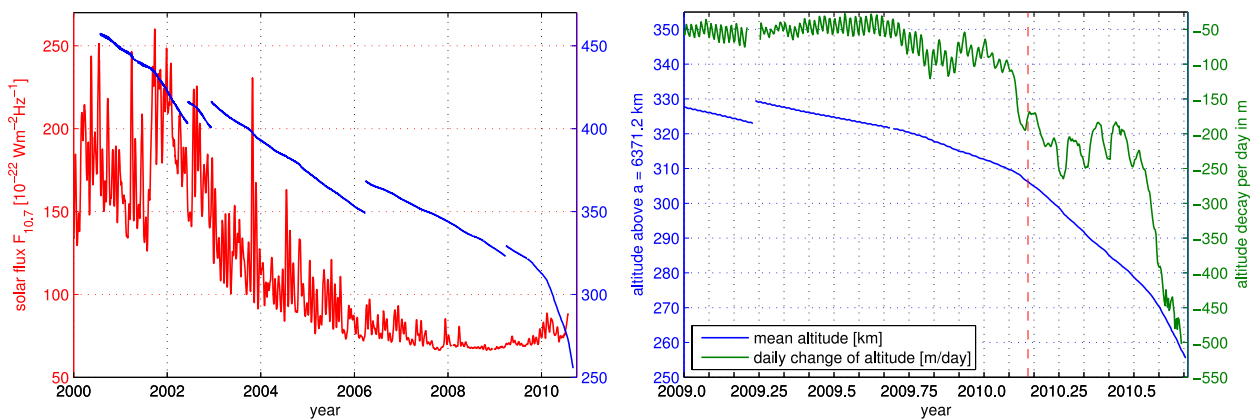
2010, 2013), the BGS models (Thomson & Lesur 2007; Thomson *et al.* 2010), and the CHAOS models (Olsen *et al.* 2006, 2009, 2010).

Until recently, the altitude of the CHAMP satellite was not sufficiently low to determine small-scale structures of the lithospheric field without along-track filtering of the data (to remove unmodelled external field contributions), and hence only terms up to spherical harmonic degree  $n = 50$  or so could be determined robustly using the comprehensive approach. Inclusion of the most recent low-altitude CHAMP data, as done here, allows determination of the lithospheric field up to at least spherical harmonic degree  $n = 85$ .

The very recent low-altitude CHAMP observations are crucial for a robust determination of small-scale lithospheric field structures. But how has CHAMP altitude evolved with time? The left-hand part of Fig. 2 shows in blue the altitude of CHAMP (with respect to a mean Earth radius of  $a = 6371.2$  km), together with the temporal evolution of the  $F_{10.7}$  solar flux (red curve). Various orbit manoeuvres are the reasons for sudden increases of the satellite altitude. Note how rising solar activity at the end of 2001 leads to a faster altitude decay, due to increased air density and thereby larger air-drag.

CHAMP altitude was about 350 km during the solar minimum years 2007–2009 and reached 300 km at the beginning of 2010. The satellite altitude for the last 2 yr of mission lifetime is shown in the right-hand part of Fig. 2, together with the mean daily change of altitude (green curve). The latter was about 50 m/d in 2009, but increased to a value of about 200 m/d during the first part of 2010, partly due to the fact that the satellite was turned by  $180^\circ$ . Before 2010 February 22 (indicated by the dashed red vertical line) CHAMP flew with its boom in the flight direction, which is a favourable condition regarding air drag, but is less optimal for attitude control. After that date CHAMP flew with the boom backward, requiring less cold gas for attitude control. After 2010 July the daily altitude decay increased rapidly to values of 500 m/d and more.

The CHAOS-4 field model presented here is the most recent version in the CHAOS model series. Previous versions are CHAOS (Olsen *et al.* 2006), xCHAOS (Olsen & Manda 2008), CHAOS-2 (Olsen *et al.* 2009) and CHAOS-3 (Olsen *et al.* 2010). In previous model versions we have concentrated on an optimal description of the core field and its temporal evolution; the static (lithospheric) field was only modelled up to relatively low spherical harmonic degrees ( $n = 60$  at most). For CHAOS-4 we extend the spherical harmonic degrees and solve for the lithospheric field up to  $n = 100$ .



**Figure 2.** Left-hand panel: solar flux index,  $F_{10.7}$  (red) and CHAMP mean altitude (blue) in dependence on time. Right-hand panel: CHAMP mean altitude (blue) and mean daily altitude decay (green) since 2009.

However, CHAMP data are not available after 2010 September since the satellite re-entered the atmosphere on 2010 September 19. In order to describe properly the time changes of the core field until mid-2013 we take advantage of the most recent scalar magnetic field observations from the Ørsted satellite and data from magnetic observatories.

CHAOS-4 is derived by merging two models, both of which are determined using the comprehensive approach but from different data sets (and using different model parameterization): spherical harmonic coefficients of the CHAOS-4 model up to degree  $n = 24$  (which includes the time-changes of the core field) are taken from model version CHAOS-4l, while coefficients describing the lithospheric field for  $n = 25$ –100 are taken from model version CHAOS-4h. The former is derived in the tradition of the earlier versions of the CHAOS series by using data between 1997.0 and 2013.5 from the three satellites Ørsted, CHAMP and SAC-C together with observatory monthly mean values, while CHAOS-4h is derived solely from the final 2 yr of low altitude CHAMP satellite data.

## 2 DATA SELECTION

### 2.1 Satellite data

We use Ørsted scalar data between 1999 March and 2013 June, Ørsted vector data between 1999 March and 2004 December, CHAMP vector and scalar data between 2000 August and 2010 September, and SAC-C scalar data between 2001 January and 2004 December. Similar data selection criteria to those chosen for determining previous versions in the CHAOS model series have been used; the main modification concerns the use of quasi-dipole (QD) latitude (Richmond 1995) instead of dipole latitude to distinguish polar from non-polar regions, and use of a revised version of the dayside merging electric field at the magnetopause when selecting data at polar latitudes and the inclusion of pre-midnight CHAMP data: for previous CHAOS versions we only used post-midnight non-polar CHAMP data to avoid contamination of the lithospheric field determination by plasma irregularities. However, these irregularities are almost absent during the solar minimum conditions of the recent years which we used to determine the lithospheric field, allowing us to include also pre-midnight data.

Data have been selected according to the following criteria: (1) we use only data from dark regions (sun  $10^\circ$  below horizon) and for which the  $RC$ -index measuring the strength of the magnetospheric ring-current does not change by more than  $2 \text{ nT hr}^{-1}$  (the determination of  $RC$  is described in Section 2.2 below); (2) all three components of the magnetic field vector are taken for QD latitudes equatorward of  $\pm 55^\circ$ , while only scalar field data (i.e. magnetic field intensity) are used for regions poleward of  $\pm 55^\circ$  or when attitude data were not available; (3) at non-polar latitudes (equatorward of  $\pm 55^\circ$  QD latitude) we require for the geomagnetic activity index  $Kp \leq 2^0$ , while for regions poleward of  $55^\circ$  QD latitude the merging electric field at the magnetopause,  $E_m$ , has to be below a certain threshold value  $E_m^{\max}$ . In previous versions of CHAOS we used a threshold value of  $E_m^{\max} = 0.8 \text{ mV m}^{-1}$  and the merging electric field as defined by Kan & Lee (1979):

$$E_m = v B_t \sin^2(\Theta/2), \quad (1)$$

where  $v$  is the solar wind speed,  $B_t = \sqrt{B_y^2 + B_z^2}$  is the magnitude of the interplanetary magnetic field in the  $y$ - $z$  plane in geocentric solar magnetospheric (GSM) coordinates [see Kivelson & Russell (1995) for a definition of the GSM coordinate system] and  $\Theta =$

$\arctan(B_y/B_z)$ . For CHAOS-4 we use a revised definition of  $E_m$ , following Newell *et al.* (2007):

$$E_m = 0.33 v^{4/3} B_t^{2/3} \sin^{8/3}(|\Theta|/2) \quad (2)$$

which gives  $E_m$  in units of  $\text{mV m}^{-1}$  if  $v$  and  $B_t$  are provided in  $\text{km s}^{-1}$  and nT, respectively. The factor 0.33 has been determined empirically in order that  $E_m$  is of about similar magnitude as the one determined using eq. (1). Solar wind speed and IMF magnetic data of temporal resolution  $\Delta t = 5 \text{ min}$ , propagated to the magnetospheric bow shock, are taken from the OMNI database ([ftp://spdf.gsfc.nasa.gov/pub/data/omni/high\\_res\\_omni/](ftp://spdf.gsfc.nasa.gov/pub/data/omni/high_res_omni/)). We consider the ‘memory effect’ of the magnetosphere-ionosphere system by using  $E_{m,24}$ , the weighted sum of  $E_m$  over the preceding 24 hr (corresponding to the preceding 24 5-min values) with weights  $w_k = \exp(-k\Delta t/0.75 \text{ h}) / \sum w_k$ . Only satellite data at polar latitudes for which  $E_{m,24} \leq 0.8 \text{ mV m}^{-1}$  have been selected.

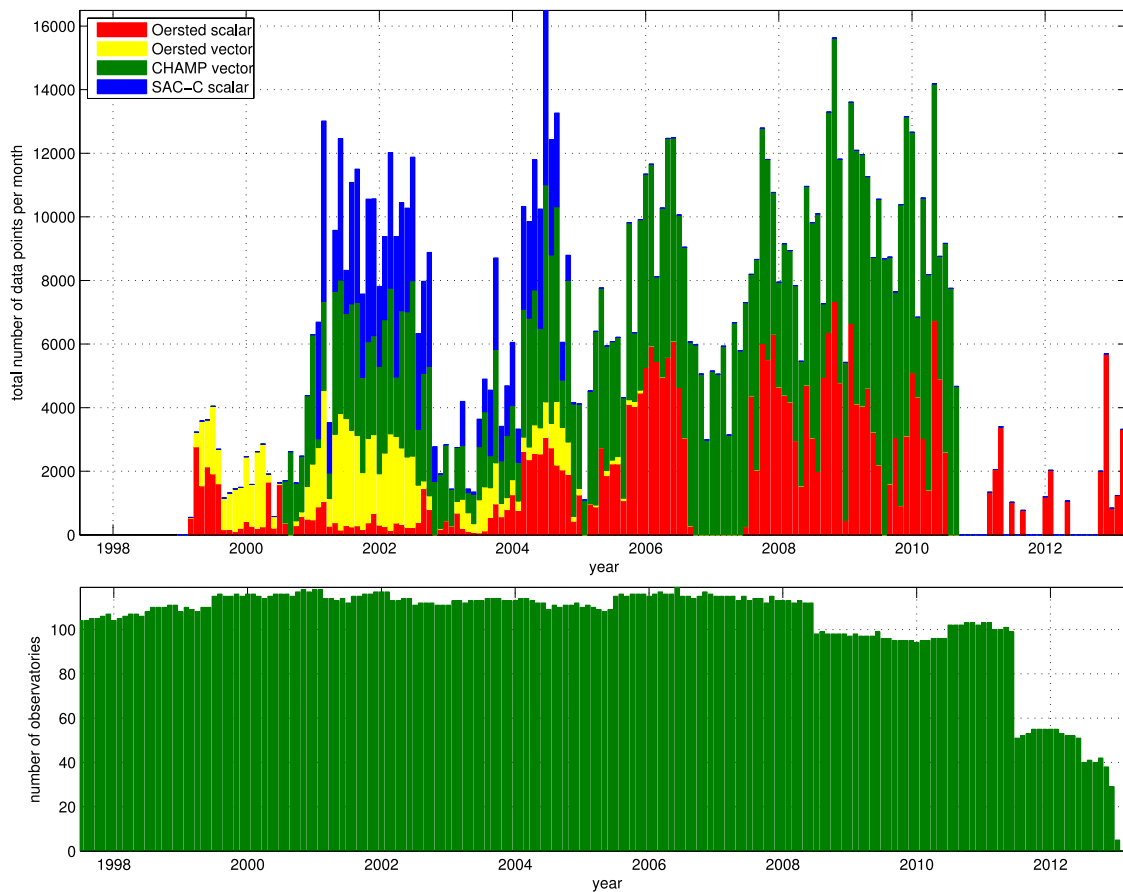
All satellite data are weighted proportional to  $\sin\theta$  (where  $\theta$  is geographic colatitude) to simulate an equal-area distribution. Anisotropic errors due to attitude uncertainty (Holme & Bloxham 1996; Holme 2000) are considered for all Ørsted vector data and for CHAMP vector data when attitude data from only one star imager were available.

In order to obtain a core field model that is up-to-date we included most recent Ørsted scalar field observations, up to 2013 June. The top panel of Fig. 3 shows the total number of non-polar magnetic satellite observations for each month. Periods with fewer data (for instance around 2003) are due to increased geomagnetic activity, while problems with the attitude stability are the reason for the Ørsted data gap around 2007 and after 2011. Beginning in 2012 November special efforts have been undertaken to improve the retrieval of Ørsted data in order to better bridge the gap between the atmospheric re-entry of CHAMP and the launch of *Swarm*. This has resulted in significantly more data since end of 2012. All Ørsted data, including the very recent ones, are freely available at <ftp.space.dtu.dk/data/magnetic-satellites/Ørsted/>.

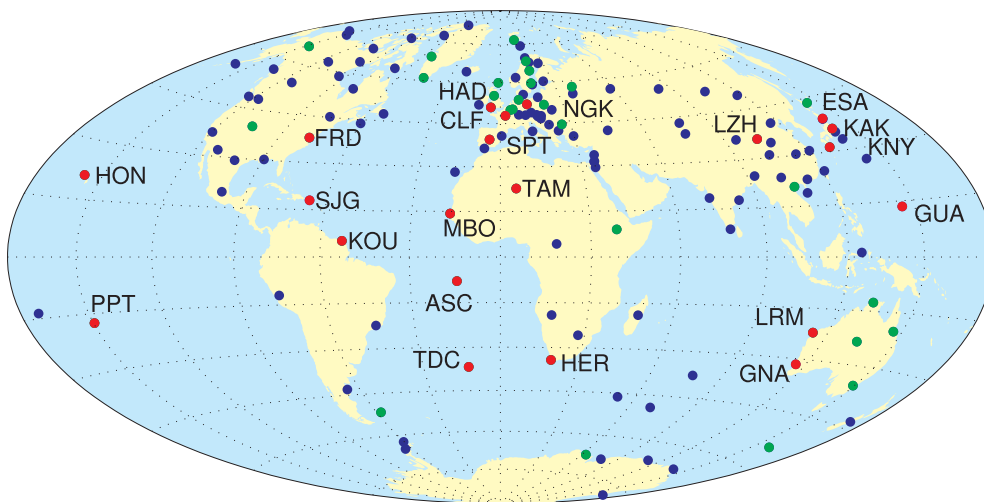
### 2.2 Ground data and the RC index describing the magnetospheric ring-current

The time-space structure of the magnetospheric field is traditionally described by the  $Dst$  index (e.g. Sugiura 1964). However, the baseline of  $Dst$  is known to change with time (e.g. Olsen *et al.* 2005; Lühr & Maus 2010), which hampers its use in geomagnetic field modelling. We therefore derived a new index, called  $RC$ , which describes the strength of the magnetospheric ring-current even during geomagnetic quiet conditions (when the baseline instabilities of  $Dst$  lead to less-optimal results).  $RC$  is derived by an hour-by-hour spherical harmonic analysis (SHA) of hourly mean values from worldwide distributed observatories at mid and low latitudes. The locations of the 21 observatories used to derive  $RC$  in this study are indicated by red dots in Fig. 4.

After removal of the core field as given by a preliminary version of CHAOS-4 from the hourly mean values, an observatory bias (accounting for the high-degree lithospheric field that is not described by that model) was subtracted from each component at each observatory. This observatory bias was determined such that the arithmetic mean value during geomagnetic quiet periods (defined as  $Kp \leq 2^0$ ,  $|dDst/dt| \leq 2 \text{ nT hr}^{-1}$ ) vanishes. Next, all data were converted from the geographic to the dipole (geomagnetic) frame. We then performed an hour-by-hour spherical harmonic analysis on the horizontal components, estimating time series of the three



**Figure 3.** Top panel: total number of non-polar satellite data (stacked histogram) as a function of time. Bottom panel: number of magnetic observatory monthly mean values as a function of time.



**Figure 4.** Location of the 146 magnetic observatories used in this study (circles). Emphasized (green or red symbols) are the 43 observatories that provided quasi-definitive (*cf.* Peltier & Chulliat 2010) data in 2013, out of which 21 observatories (red dots with three letter code given) were used for calculating the *RC* index.

spherical harmonic expansion coefficients  $v_1^m$  of degree  $n = 1$  and order  $m = 0$  and 1. For this spherical harmonic analysis we used only those observatories that were (for the UT hour in consideration) located in the night-side sector (local time between 18 and 06). Due to this selection only between 5 (for UT = 08) and 16 (for UT = 20) out of the 21 observatories have been used, with the number of observatories changing from hour to hour.

The *RC* index is defined as  $RC = -v_1^0$ ; the minus sign is introduced in order to make *RC* compatible to the definition of *Dst* as the southward component of the magnetic field at the dipole equator.

Since *RC* is derived from the horizontal components only, it consists of the sum of magnetospheric and induced part, similar to *Dst* [this is the reason why we denote the spherical harmonic expansion coefficients as  $v_1^m$ ; it is the sum of the external (magnetospheric)

coefficient  $q_1^m$  and the internal (induced) coefficient  $g_1^m$ ]. In a final step we decompose  $RC(t) = \epsilon(t) + \iota(t)$  into external and induced parts, respectively, using a 1-D model of electrical conductivity of the mantle as described in Olsen *et al.* (2005).

To stabilize the determination of the CHAOS-4 core field time changes and to extend it back in time before February 1999 (the launch of the Ørsted satellite) we supplement the satellite data with annual differences of revised observatory monthly means of the north, east and vertical downward components ( $X, Y, Z$ ) for the time interval 1997.5 to 2013.5. These monthly means have been derived from hourly mean values of 146 observatories (shown in Fig. 4) which have been carefully checked for trends, spikes and other errors as described in Macmillan & Olsen (2013). We removed from the observations (i) synthetic model values of the ionospheric (plus induced) field as predicted by the CM4 model (Sabaka *et al.* 2004), which takes as input 3-monthly means of  $F_{10.7}$  solar flux and (ii) model values for the magnetospheric (plus induced) field as predicted by a preliminary version of CHAOS-4, parametrized by the  $RC$  index and its decomposition into external and induced parts. After subtracting these field corrections from the observed hourly mean values we calculate the robust mean [using Huber weights with a tuning constant of 1.5, cf. Huber (2011)] of all hourly mean values of a given month, for each of the 146 observatories and each of the three elements  $X, Y$  and  $Z$ . Finally, we take the annual differences of the resulting revised monthly means (annual difference value at time  $t$  is obtained by taking the difference between those at  $t + 6$  months and  $t - 6$  months to eliminate a remaining annual variation in the data). This yields 19 299 values of the first time derivative of the vector components,  $dX/dt, dY/dt, dZ/dt$ , for 146 observatories. Their distribution in time is given in the bottom panel of Fig. 3.

### 3 MODEL PARAMETRIZATION

Parametrization of the CHAOS-4 model follows closely that of the previous versions in the CHAOS model series. The model consists of spherical harmonic expansion coefficients describing the magnetic field vector in an Earth-centred earth-fixed (ECEF) coordinate system and sets of Euler angles needed to rotate the satellite vector readings from the magnetometer frame to the star imager frame. The magnetic field vector in the ECEF frame,  $\mathbf{B} = -\nabla V$ , is derived from a magnetic scalar potential  $V = V^{\text{int}} + V^{\text{ext}}$  consisting of a part,  $V^{\text{int}}$ , describing internal (core and lithospheric) sources, and a part,  $V^{\text{ext}}$ , describing external (mainly magnetospheric) sources and their Earth-induced counterparts. Both parts are expanded in terms of spherical harmonics.

For the internal part this yields

$$V^{\text{int}} = a \sum_{n=1}^{N_{\text{int}}} \sum_{m=0}^n (g_n^m \cos m\phi + h_n^m \sin m\phi) \left(\frac{a}{r}\right)^{n+1} P_n^m(\cos\theta), \quad (3)$$

where  $a = 6371.2$  km is a reference radius,  $(r, \theta, \phi)$  are geographic coordinates,  $P_n^m$  are the associated Schmidt semi-normalized Legendre functions,  $\{g_n^m, h_n^m\}$  are the Gauss coefficients describing internal sources, and  $N_{\text{int}}$  is the maximum degree and order of the internal expansion.

As mentioned above, the final CHAOS-4 model is found by merging two submodels, called CHAOS-4l, resp. CHAOS-4h. They differ in the maximum spherical harmonic degree  $N_{\text{int}}$  of the static field, in the temporal parametrization of the low-degree (core field) terms, and in the data sets that have been used to derive these submodels. Details are given below in Sections 3.1 and 3.2.

Common to both submodels is the parametrization of external fields, with an expansion of the near magnetospheric sources (magnetospheric ring current) in the solar magnetic (SM) coordinate system (up to  $n = 2$ , with special treatment of the  $n = 1$  terms) and of remote magnetospheric sources (e.g. magnetotail and magnetopause currents) in geocentric solar magnetospheric (GSM) coordinates (also up to  $n = 2$ , but restricted to order  $m = 0$ ):

$$V^{\text{ext}} = a \sum_{n=1}^2 \sum_{m=0}^n (q_n^m \cos mT_d + s_n^m \sin mT_d) \left(\frac{r}{a}\right)^n P_n^m(\cos\theta_d) + a \sum_{n=1}^2 q_n^{0,\text{GSM}} R_n^0(r, \theta, \phi), \quad (4)$$

where  $\theta_d$  and  $T_d$  are dipole colatitude and dipole local time, respectively. The degree-1 coefficients in SM coordinates depend explicitly on time and are further expanded as

$$q_1^0(t) = \hat{q}_1^0 \left[ \epsilon(t) + \iota(t) \left(\frac{a}{r}\right)^3 \right] + \Delta q_1^0(t) \quad (5a)$$

$$q_1^1(t) = \hat{q}_1^1 \left[ \epsilon(t) + \iota(t) \left(\frac{a}{r}\right)^3 \right] + \Delta q_1^1(t) \quad (5b)$$

$$s_1^1(t) = \hat{s}_1^1 \left[ \epsilon(t) + \iota(t) \left(\frac{a}{r}\right)^3 \right] + \Delta s_1^1(t), \quad (5c)$$

where the term in brackets describes the magnetic field contribution due to the magnetospheric ring-current and its Earth-induced counterpart as given by the RC index (see Section 2.2),  $RC(t) = \epsilon(t) + \iota(t)$ . If  $RC$  provides a perfect description of the magnetospheric field at satellite altitude the values of the regression coefficients would be  $\hat{q}_1^0 = -1, \hat{q}_1^1 = \hat{s}_1^1 = 0$  with vanishing ‘RC baseline corrections’  $\Delta q_1^0, \Delta q_1^1$  and  $\Delta s_1^1$ . We allow for deviations from these values and co-estimate the regression factors  $\hat{q}_1^0, \hat{q}_1^1, \hat{s}_1^1$  and solve for the baseline corrections in bins of 5 d (for  $\Delta q_1^0$ ) and 30 d (for  $\Delta q_1^1, \Delta s_1^1$ ), respectively. Note that these bin-sizes are considerably longer than for previous versions of the CHAOS model series (with bin-sizes of 12 hr and 5 d, respectively) since the baseline of  $RC$  is more stable compared to that of  $Dst$ , allowing for longer bin-sizes.

In the following we describe in more detail the data and model parametrization used for the two submodels CHAOS-4l and CHAOS-4h.

#### 3.1 CHAOS-4l

CHAOS-4l (where the ‘l’ stands for ‘low degree’) is determined using the whole data set (Ørsted, CHAMP and SAC-C satellite data plus observatory monthly means) described above, with a sampling rate of the satellite data of 60 s.

The maximum spherical harmonic degree of the internal field,  $N_{\text{int}} = 80$ , is higher than that of CHAOS-2 and CHAOS-3. Internal Gauss coefficients  $\{g_n^m(t), h_n^m(t)\}$  up to  $n = 20$  are time-dependent; this dependence is described by order 6 B-splines (Schumaker 1981; De Boor 2001) with a 6-month knot separation and fivefold knots at the endpoints  $t = 1997.0$  and  $t = 2013.5$ . This yields 32 interior knots (at 1997.5, 1998.0, ..., 2013.0) and six exterior knots at each endpoint, 1997.0 and 2013.5, resulting in 38 basic B-spline functions  $M_l(t)$ . Internal Gauss coefficients for degrees  $n = 21-80$  are static. Time-dependent terms (for degrees  $n = 1-20$ ) and static terms (for  $n = 21-80$ ) together result in a total of 22 840 internal Gauss coefficients.

The total number of external field parameters is 1230, which is the sum of five SM terms ( $q_2^m, s_2^m$  for  $m = 0 - 2$ ), 3 RC

regression coefficients  $\tilde{q}_1^0, \tilde{q}_1^1, \tilde{s}_1^1$ , two GSM coefficients ( $q_n^{1,\text{GSM}}, q_n^{2,\text{GSM}}$ ), 900 baseline corrections  $\Delta q_1^0$  and  $2 \times 160$  baseline corrections  $\Delta q_1^1, \Delta s_1^1$ .

As part of the model estimation we solve for the Euler angles describing the rotation between the vector magnetometer frame and the star sensor frame. For Ørsted this yields two sets of Euler angles (one for the period before 2000 January 24 when the onboard software of the star-imager was updated and one for the period after that date), while for CHAMP we solve for Euler angles in bins of 10 d (i.e. 367 sets of angles). This yields additional  $3 \times (2 + 367) = 1107$  model parameters. Together with the 22 840 parameters describing the internal field and the 1230 model parameters describing the external field this yields a total number of  $22\,840 + 1230 + 1107 = 25\,177$  model parameters. These model parameters are estimated from 727 776 scalar data and  $3 \times 651\,397$  vector data by means of a regularized iteratively reweighted least-squares approach, using Huber weights, minimizing the cost function

$$\mathbf{e}^T \underline{\underline{C}}^{-1} \mathbf{e} + \lambda_3 \mathbf{m}^T \underline{\underline{\Lambda}}_3 \mathbf{m} + \lambda_2 \mathbf{m}^T \underline{\underline{\Lambda}}_2 \mathbf{m}, \quad (6)$$

where  $\mathbf{m}$  is the model vector, the residuals vector  $\mathbf{e} = \mathbf{d}_{\text{obs}} - \mathbf{d}_{\text{mod}}$  is the difference between observation  $\mathbf{d}_{\text{obs}}$  and model prediction  $\mathbf{d}_{\text{mod}}$ , and  $\underline{\underline{C}}$  is the data covariance matrix.

$\underline{\underline{\Lambda}}_3$  and  $\underline{\underline{\Lambda}}_2$  are block diagonal regularization matrices which constrain the third, respectively second, time derivatives of the core field. Contrary to previous versions in the CHAOS model series (for which field intensity  $|\mathbf{B}|$  is regularized) we constrain the radial field component  $B_r$  at the core surface in CHAOS-4l.  $\underline{\underline{\Lambda}}_3$  minimizes the mean squared magnitude of  $|\frac{\partial^3 B_r}{\partial t^3}|$ , integrated over the core surface  $\Omega_c$  (radius  $c = 3485$  km) and averaged over time:

$$\left\langle \left| \frac{\partial^3 B_r}{\partial t^3} \right|^2 \right\rangle = \frac{1}{\Delta t} \int_{t=1997}^{2013.5} \int \left| \frac{\partial^3 B_r}{\partial t^3} \right|^2 d\Omega_c dt = \mathbf{m}^T \underline{\underline{\Lambda}}_3 \mathbf{m} \quad (7)$$

with  $\Delta t = 2013.5 - 1997.0 = 16.5$  yr.

Regularization of the third time derivative alone leads to unconstrained field oscillations. To avoid this we also minimize  $|\frac{\partial^2 B_r}{\partial t^2}|$  at the core surface at the model endpoints  $t = 1997.0$  and  $2013.5$ . This is implemented via the regularization matrix  $\underline{\underline{\Lambda}}_2$ . Note that  $\underline{\underline{\Lambda}}_2$  only acts on 12 (the first and last six) of the 38 spline basis functions.

The two parameters  $\lambda_3$  and  $\lambda_2$  control the strength of the regularizations. We considered several values for these parameters and finally selected  $\lambda_2 = 10$  (nT yr<sup>-2</sup>)<sup>-2</sup> and  $\lambda_3 = 0.33$  (nT yr<sup>-3</sup>)<sup>-2</sup>. The axial dipole coefficient  $g_0^0$  is treated separately, with  $\lambda_3$  increased to  $10$  (nT yr<sup>-3</sup>)<sup>-2</sup>, since it is the internal field coefficient most affected by unmodelled external field fluctuations.

### 3.2 CHAOS-4h

CHAOS-4h is the high-degree part of the CHAOS-4 model (as indicated by the suffix ‘h’). The maximum spherical harmonic degree of the internal model part is  $N_{\text{int}} = 100$ . CHAOS-4h is derived solely from CHAMP data obtained between 2008 September and 2010 September. The data sampling interval is 30 s. A quadratic time dependence is used to describe the temporal variation of core field coefficients up to degree  $n = 16$ ; coefficients between  $n = 17$ –100 are assumed to be static. This results in a total of 10 776 internal Gauss coefficients. The number of model parameters describing the external field is 1579. Euler angles of the transformation between the magnetometer and the star sensor frame are solved for in bins of 10 d, resulting in  $3 \times 72 = 216$  additional model parameters. In

total model CHAOS-4h consists of 12 571 model parameters which are estimated from 118 184 scalar data and  $3 \times 313\,840$  vector data.

No temporal regularization is applied, but for spherical harmonic degrees above  $n = 85$  we regularize the mean squared value of  $|B_r|$  over the Earth’s surface (i.e. at about the depth of the crustal field source). Similar to eq. (6) we minimize the cost function

$$\mathbf{e}^T \underline{\underline{C}}^{-1} \mathbf{e} + \lambda_0 \mathbf{m}^T \underline{\underline{\Lambda}}_0 \mathbf{m}, \quad (8)$$

using iteratively reweighted least-squares with Huber weights, where  $\underline{\underline{\Lambda}}_0$  constrains the mean squared magnitude of  $|B_r|$  for  $n > 85$ , integrated over the Earth surface  $\Omega_a$  (radius  $r = a$ ):

$$\langle |B_r|^2 \rangle = \int |B_r|^2 d\Omega_a dt = \mathbf{m}^T \underline{\underline{\Lambda}}_0 \mathbf{m}. \quad (9)$$

$\underline{\underline{\Lambda}}_0$  is a diagonal matrix with elements  $(n+1)^2/(2n+1)$  for  $n > 95$ , zero elements for  $n < 85$ , and a smooth transition between  $n = 85$  and 95. A threshold of  $n = 85$  was chosen since preliminary, not-regularized, models indicate that the spectra tend to increase rapidly for  $n > 85$ . We have chosen a damping parameter of  $\lambda_0 = 1$  nT<sup>-2</sup>.

When estimating high-degree magnetic field models using near-polar orbiting satellites, special attention has to be paid to the polar gaps, which are the regions around the geographic poles of half-angle  $|90^\circ - i|$  where  $i$  is the inclination of the orbit. Inclination of the CHAMP orbit is  $i = 87.3^\circ$ , which results in a polar gap of half-angle  $2.7^\circ$ . Such a gap is especially problematic for the proper determination of zonal coefficients  $g_n^0$  above  $n = 60$ . To avoid ringing at the poles we therefore add 1800 synthetic field values of  $B_r$ , synthesized within the polar gap from model CHAOS-4l up to degree  $n = 60$ . This suppresses the zonal coefficients  $g_n^0$  for  $n > 60$ .

### 3.3 Merging the two submodels CHAOS-4l and CHAOS-4h

The final CHAOS-4 model is obtained by combining the spherical harmonic coefficients of the two submodels CHAOS-4l and CHAOS-4h. Spherical harmonic coefficients up to  $n = 24$  (which includes the time-changing core field) and the external field are taken from CHAOS-4l, while the lithospheric field coefficients for degrees  $n = 25$ –100 are taken from CHAOS-4h. The transition occurs at degree  $n = 25$  where the degree correlation  $\rho_n$  between the two models reaches a maximum of  $\rho_n = 0.999$  while their relative difference (degree variance of model difference divided by degree variance of CHAOS-4) is less than 0.2 per cent. For all degrees up to  $n = 50$  the correlation  $\rho_n > 0.996$  and the relative variance difference is below 1 per cent.

## 4 RESULTS AND DISCUSSION

The total number of data points, residual means and rms values of the two submodels are listed in Table 1. Means and rms are the weighted values calculated from the model residuals  $\mathbf{e} = \mathbf{d}_{\text{obs}} - \mathbf{d}_{\text{mod}}$  using the Huber weights obtained in the final iteration. The statistics of model version CHAOS-4h, with rms values of about 1.8 nT for the scalar field and of 2.2–2.4 nT for the two remaining vector components, impressively demonstrate the capability of CHAMP. The rms values of model version CHAOS-4l are somewhat larger, which is likely due to the longer time span of that model (14 yr compared to only 2 yr for CHAOS-4h), the probably less optimal data quality at the beginning of the CHAMP mission, and the fact that data from four different sources (Ørsted, CHAMP, SAC-C and

**Table 1.** Number  $N$  of data points, mean and rms misfit (in nT for the satellite data, and in nT yr<sup>-1</sup> for the observatory data) for CHAOS-4l and CHAOS-4h. Statistics for the vector components are given in the  $(B_r, B_\theta, B_\phi)$  coordinate system and in a coordinate system  $(B_2, B_\perp, B_3)$  defined by the bore-sight of the star imager and the ambient field direction (cf. Olsen *et al.* 2000, for details).

			CHAOS-4l			CHAOS-4h		
Component			$N$	Mean	Rms	$N$	Mean	Rms
Satellite	Ørsted	$F_{\text{polar}}$	122,363	0.37	3.35			
		$B_r$	89,833	-0.14	5.13			
		$B_\theta$	89,833	-0.24	5.23			
		$B_\phi$	89,833	-0.05	4.60			
		$F_{\text{nonpolar}} + B_B$	389,229	0.13	2.35			
	CHAMP	$B_\perp$	89,833	-0.04	7.40			
		$B_3$	89,833	0.12	3.40			
		$F_{\text{polar}}$	190,578	-0.30	4.85	118,184	0.02	4.16
		$B_r$	542,265	0.03	2.56	313,840	0.00	1.83
		$B_\theta$	542,265	0.05	3.58	313,840	0.05	2.46
SAC-C	$B_\phi$	542,265	0.14	2.95	313,840	0.00	2.11	
	$F_{\text{nonpolar}} + B_B$	542,265	-0.11	2.10	313,840	0.00	1.68	
	$B_\perp$	542,265	-0.02	3.36	313,840	-0.02	2.27	
	$B_3$	542,265	0.06	3.47	313,840	0.07	2.51	
	$F_{\text{polar}}$	26,051	0.54	3.81				
Observatory	$F_{\text{nonpolar}}$	89,388	0.41	2.69				
	$dB_r/dt$	19,299	0.11	3.78				
	$dB_\theta/dt$	19,299	-0.13	3.80				
		$dB_\phi/dt$	19,299	-0.02	2.98			

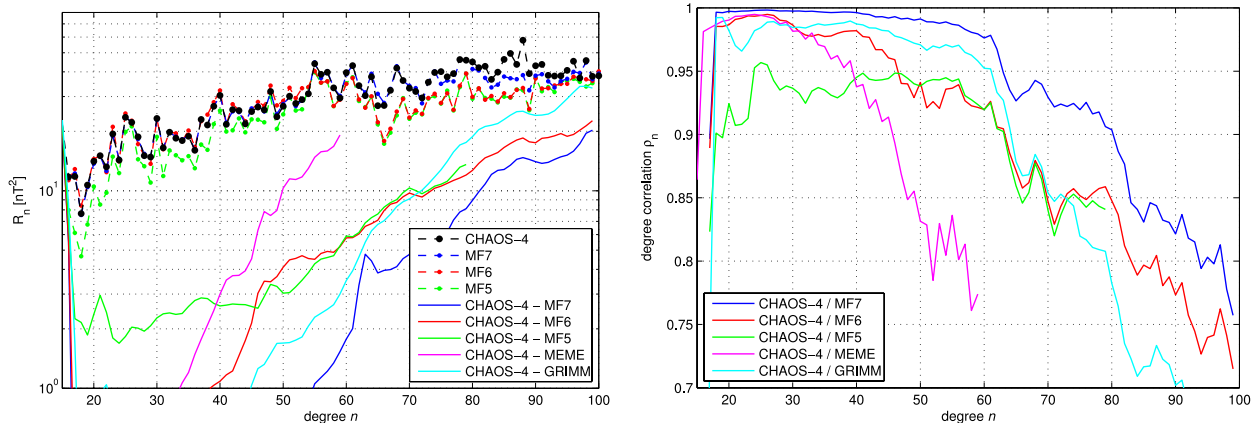
magnetic observatories) have been combined. The reason for the larger CHAMP residuals compared to the two other satellites is likely due to its lower altitude, which means it is closer to the disturbing ionosphere as well as to unmodelled signals from the lithosphere.

#### 4.1 The lithospheric field

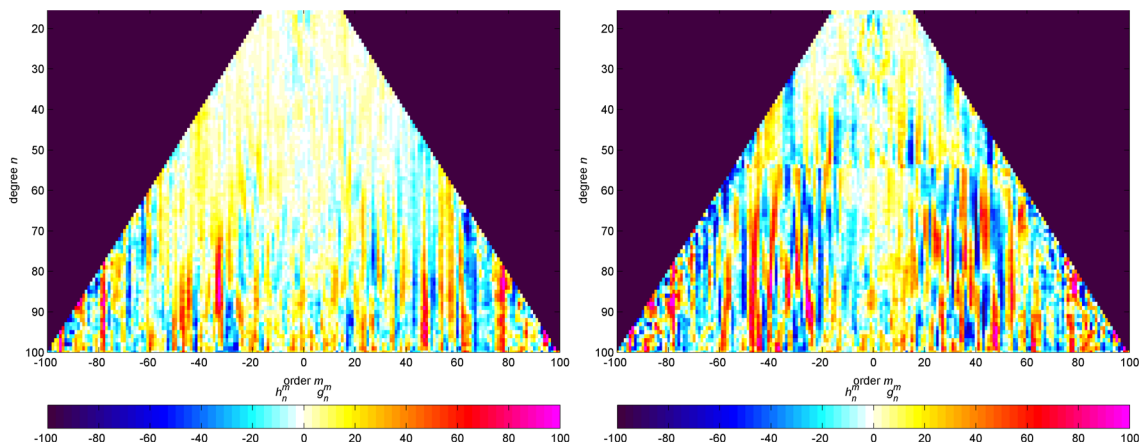
The left-hand part of Fig. 5 shows the spatial power spectrum of CHAOS-4 (black dots) and of various versions of the MF model series determined by Stefan Maus and co-workers. Up to spherical harmonic degree  $n = 80$  the power of CHAOS-4 is very similar to that of MF7 (Maus 2010), which is the most recent version in the MF model series. The power of previous versions in the MF series is smaller than that of MF7 and CHAOS-4. Noteworthy is how the spectral power of the MF models increases with increasing version number: for degrees  $n = 16$ –60 the power of MF5 is 15 per cent below that of CHAOS-4 while for the later versions MF6 and MF7 it

is similar to that of CHAOS-4 within 1 per cent. For higher degrees ( $n = 61$ –80) the power of MF6 is as much as 30 per cent below that of CHAOS-4. The power of the lithospheric signal in MF7 has increased compared to that in MF6 and is very close to that of CHAOS-4. This change in power illustrates the effect of high-pass filtering that was applied in the determination of the MF models; earlier model versions have been derived using heavier filtering compared to more recent versions, resulting in an underestimation of the lithospheric signal. Thébault *et al.* (2012) discuss the pros and cons of an along-track filtering on the determination of the lithospheric field.

Filtering has been relaxed in the most recent versions of the MF series, which clearly results in a stronger lithospheric field. For MF7 filtering is only applied for degrees above  $n = 77$ ; up to that degree the lithospheric field of CHAOS-4 and MF7 are very similar. But following the argument of the above results, it seems possible that MF7 may underestimate the lithospheric signal for degrees above  $n = 77$ , when filtering is again applied. The



**Figure 5.** Left-hand panel: power spectra of the static field (dots) and of the field differences (solid lines). Right-hand panel: degree correlation between various model pairs.



**Figure 6.** Sensitivity matrix (normalized coefficient differences in per cent) between CHAOS-4 and MF7 (left-hand panel), resp. MF6 (right-hand panel).

*Swarm* constellation mission consisting of three identical satellites, two of which are flying side-by-side at low altitude, will give an excellent opportunity to investigate this further. Experiments based on *Swarm*-type synthetic data indicate that the lithospheric field up to degree  $n = 150$  can be robustly determined without along-track filtering (Sabaka & Olsen 2006; Olsen *et al.* 2007; Tøffner-Clausen *et al.* 2010; Sabaka *et al.* 2013).

In addition to the spectra of the various lithospheric field signals, the left-hand part of Fig. 5 shows the spectra of the differences between CHAOS-4 and various other field models. The difference is smallest when comparing CHAOS-4 with MF7 which gives an indication of the present uncertainty in lithospheric field modelling. The fact that the lithospheric field power is well above the power of the difference between CHAOS-4 and MF7—two models which have been derived using rather different approaches—confirms that lithospheric field structures up to at least degree  $n = 80$ , corresponding to a horizontal wavelength of 500 km, are currently robustly determined.

The right-hand part of Fig. 5 shows the degree correlation  $\rho_n$  [see eq. 4.23 of Langel & Hinze (1998) for a definition] between CHAOS-4 and various other field models. The highest correlation is obtained with MF7, where the degree correlation is above 0.97 for  $n \leq 60$  and above 0.85 for  $n \leq 80$ .

The left-hand part of Fig. 6 shows the sensitivity matrix  $S(n, m)$ , which is the relative difference between each coefficient of CHAOS-4 and MF7 in a degree versus order matrix; the right-hand part shows  $S(n, m)$  of the difference between CHAOS-4 and MF6. This figure confirms the better agreement between CHAOS-4 and MF7 compared to MF6. There exist, however, certain spherical harmonic orders for which the difference between CHAOS-4 and MF7 is especially large, for instance around  $m = 60$ . The reason for these differences is unknown.

A map of the radial magnetic field at Earth's surface, calculated from coefficients between degrees  $n = 16$  and 85, is shown in the upper part of Fig. 7. The bottom part shows the difference in  $B_r$  between CHAOS-4 and MF7. The largest differences occur in polar regions: Maximum field difference is 55 nT (75 nT) at Northern (Southern) polar latitudes while maximum difference at non-polar latitudes is only 33 nT. This indicates the challenge of proper lithospheric field determination in the vicinity of the ever present magnetic external field perturbations at high latitudes. The reason for the north–south oriented stripes at non-polar latitudes is not clear; it could be caused by contamination from unmodelled magnetospheric contributions in CHAOS-4

or due to the along-track filtering used during the derivation of MF7.

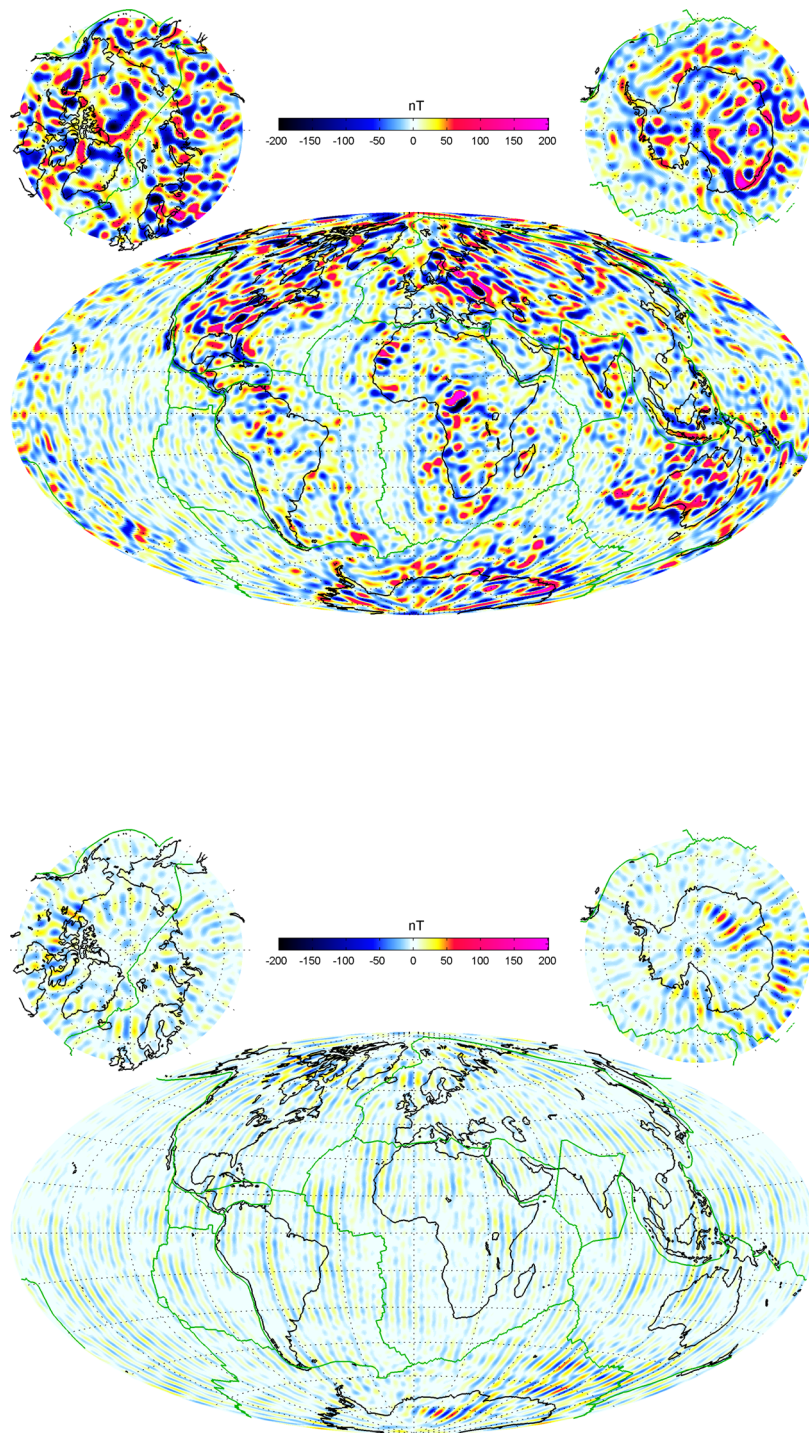
#### 4.2 The core field, secular variation and secular acceleration

An important question is how well we are able to model the time changes in the core field, especially after 2010 September when only the Ørsted satellite and observatory monthly means are available. Fig. 8 shows the secular variation (SV) and secular acceleration (SA) spectra for various epochs. The colour of the spectrum indicates the model epoch, varying from blue (for  $t_0 = 2000.0$ ) to red (for  $t_0 = 2012.0$ ). We find similar power spectra for the SV captured in CHAOS-4 before and after 2010 September, with no obvious discontinuity. There is considerable time-dependence in the SA spectra below degree 8, also after 2010, while above degree 12 the form of the spectra is essentially controlled by the temporal regularization.

Fig. 9 presents annual differences of revised monthly means, in dipole co-ordinates, from the observatories Kakioka – KAK (Japan), MBour - MBO (Senegal) and Canberra – CNB (Australia) together with SV predictions from both the CHAOS-4 model and an earlier version CHAOS-4 $\alpha$ , derived in 2010 December. Both models do an excellent job of fitting the observations, particularly when the scatter is low (e.g. KAK  $dB_r/dt$ ,  $dB_\phi/dt$ ; MBO,  $dB_r/dt$ ; CNB  $dB_r/dt$ ). The extrapolation of CHAOS-4 $\alpha$  to beyond 2010 September notably fails (e.g. MBO  $dB_\theta/dt$ , CNB  $dB_\theta/dt$ ) while CHAOS-4 does better, illustrating its reliability in the post-CHAMP era. CHAOS-4 also does a much better job at fitting rapid changes in the Atlantic region that have occurred since 2009 (e.g. in  $dB_r/dt$  at MBO, also in  $dB_r/dt$  at Ascension Island and in Hermanus, South Africa not shown). Localized, rapid, core field changes, particularly at mid-to-low latitudes have previously been studied by Olsen & Mandea (2008) and Chulliat *et al.* (2010); the latest CHAOS-4 model is evidently suitable for further study of such events, since it succeeds in capturing the amplitude and phase of the latest rapid field changes, especially in the low latitude Atlantic sector.

Fig. 10 presents a map of the radial component of the SV in 2008, truncated at spherical harmonic degree  $n = 15$  and plotted at the core surface. As previously pointed out by Holme *et al.* (2011), the SV is noticeably low under the Pacific, with the strongest SV generally occurring under the Atlantic hemisphere. Finlay *et al.*



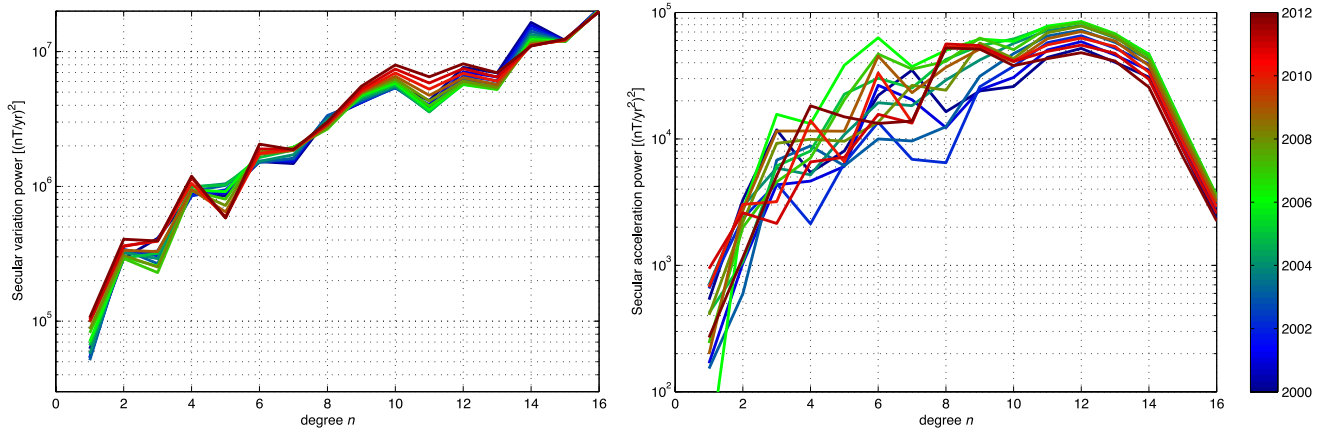


**Figure 7.** Top panel: map of the radial lithospheric magnetic field at ground, calculated from coefficients of degrees  $n = 16$ – $85$ . Bottom panel: radial field differences between CHAOS-4 and MF7.

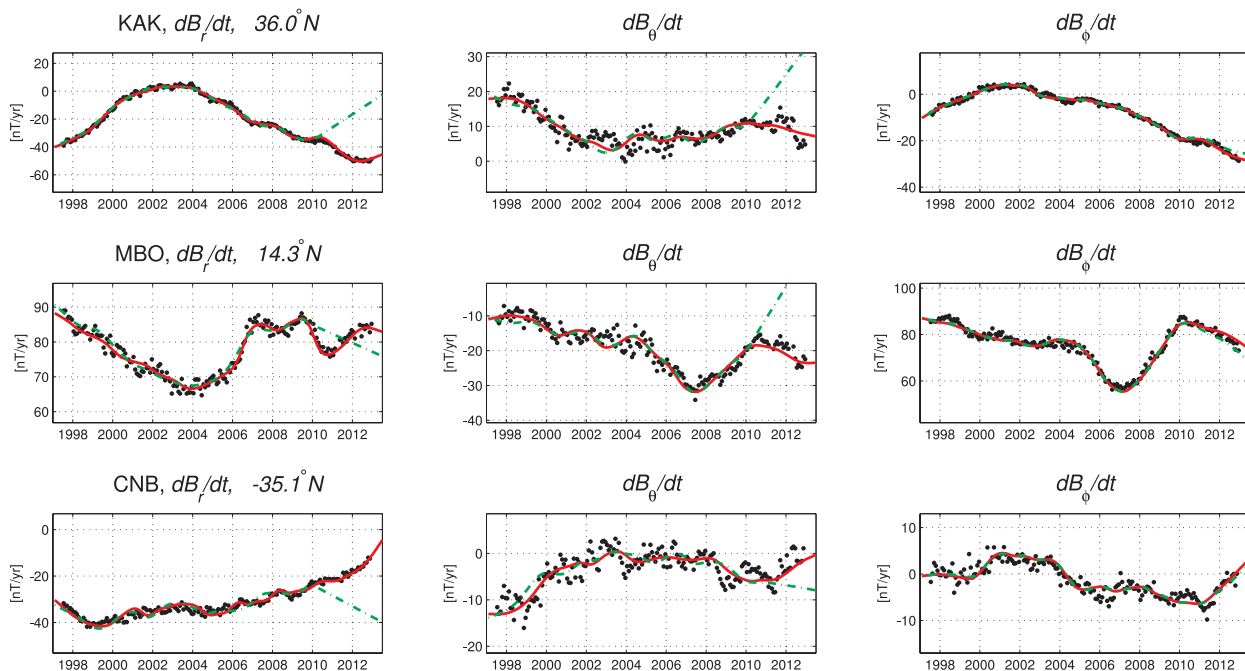
(2012) proposed that enhanced SV under the Atlantic was due to a planetary scale quasi-geostrophic gyre seen in this region in frozen flux core flow inversions (Pais & Jault 2008; Gillet *et al.* 2009). Aubert *et al.* (2013) have recently presented a new, self-consistent, geodynamo model that demonstrates how the coupled dynamics of the inner core, outer core and mantle system can produce this gyre and the general concentration of SV at mid-to-low latitudes under the Atlantic hemisphere as seen in Fig. 10.

Looking at the SV at the core surface in the polar regions, Fig. 10 provides further evidence that the SV is very low in the Southern

polar regions (see also Holme *et al.* 2011), while in the northern polar region there is a pair of intense, large scale, SV patches. These are related to the present westward motion of a prominent flux patch located under Siberia (Blokhin & Gubbins 1985) and the evolution of a similar patch under Canada/Alaska. Although there are similar high latitude flux patches in the Southern hemisphere, these do not change in such a consistent fashion, hence they generate little SV. This difference in SV between the northern and southern hemisphere polar regions is likely related to different core flows within the northern and southern tangent cylinders in the outer core. The



**Figure 8.** Mauersberger-Lowes spectra of first (left-hand panel) and second (right-hand panel) time derivatives at the core surface. The different curves correspond to different epochs  $t_0$  indicated by the colour, from  $t_0 = 2000.0$  (blue) in steps of 1 yr until  $t_0 = 2012.0$  (red).

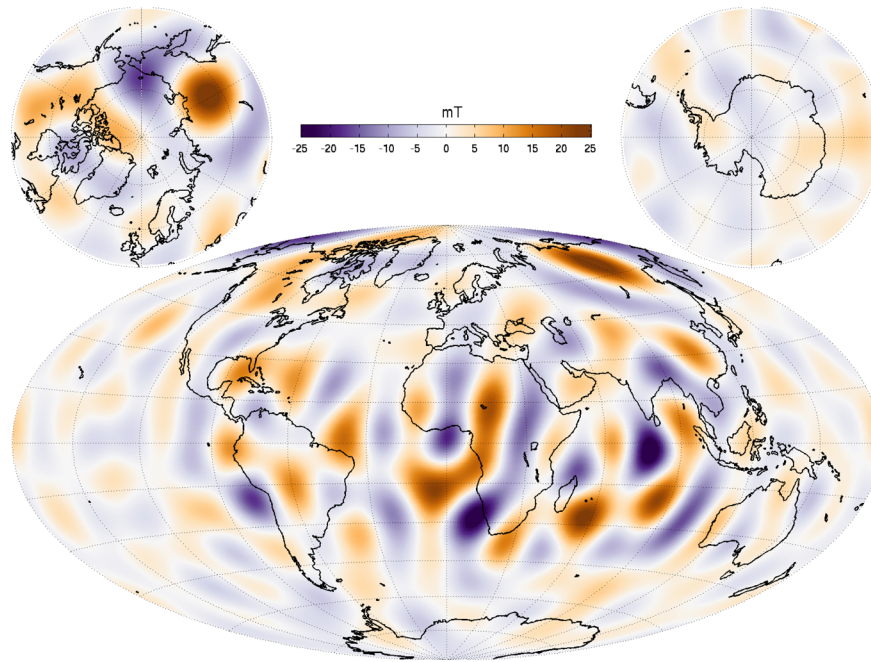


**Figure 9.** Annual differences of revised magnetic observatory monthly means (black dots) compared to predictions from CHAOS-4 (red line) and from the earlier CHAOS-4 $\alpha$  (green dashed line). Top row is Kakioka, Japan, middle row is MBour, Senegal, bottom row is Canberra, Australia.

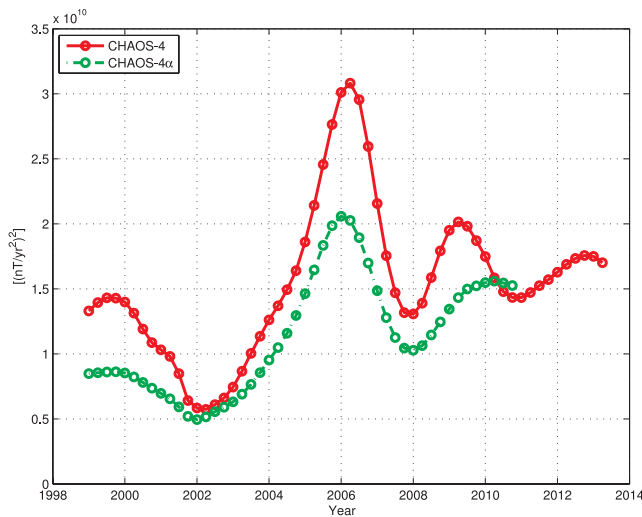
tangent cylinder acts as a dynamic barrier due the Proudman-Taylor theorem, giving rise to different, essentially decoupled, dynamics inside the two polar regions (Olson & Aurnou 1999; Sreenivasan & Jones 2005). Our results suggest there may be westward flow within the northern tangent cylinder but not within the southern tangent cylinder.

Next, we turn to the second time derivative of the core field, also known as the secular acceleration or SA. Chulliat *et al.* (2010) have pointed out that there was a large peak in the SA at the core surface in 2006. They showed how this pulse event effectively produced jerk signatures in the field at Earth's surface in 2003 and 2007. In Fig. 11, we plot the SA power at the core surface [following Chulliat *et al.* (2010)], summed over spherical harmonics up to degree  $n = 8$ , for which the time-dependence in the SA is well determined. In CHAOS-4 we find not only a dominant peak in 2006, as was the case in earlier members of the CHAOS model series, but also a secondary

maximum in 2009. This second peak was not distinguishable from end effects in our earlier model CHAOS-4 $\alpha$ , but is now seen to be a distinct signal. Following the logic of Chulliat *et al.* (2010) and Chulliat & Maus (2013) this pulse should produce jerk-like behaviour in 2008/2009 and 2011, especially close to where the pulse is strongest, at low latitudes in the mid-Atlantic sector—this is compatible with the rapid SV observed at these times in ground observatories such as MBour and Ascension Island. The amplitude of the SA pulse in 2009, needed to fit the satellite and ground observatory data, is clearly less than that in 2006. Note there are good quality CHAMP data available for a year either side of both 2006 and 2009. The core processes responsible for these SA pulses are not yet certain, but the 2009 feature, following close after the 2006 feature but with lower amplitude, could conceivably be some form of a damped oscillation response to a primary perturbation in 2006.



**Figure 10.** Secular variation of the radial magnetic field component,  $dB_r/dt$ , up to degree  $n = 15$ . Plotted at the core surface for 2008.0.



**Figure 11.** Time variation of the squared field acceleration power (up to degree 8) at the core surface.

**Table 2.** Model coefficients of the magnetospheric part of CHAOS-4. For details see Section 4.3.

$\Delta q_1^0$	=	1.90 nT	$\hat{q}_1^0$	=	-0.91
$\Delta q_1^1$	=	-1.05 nT	$\hat{q}_1^1$	=	-0.01
$\Delta s_1^1$	=	-0.17 nT	$\hat{s}_1^1$	=	-0.01
$q_2^0$	=	-1.06 nT	$q_1^{0,\text{GSM}}$	=	11.89 nT
$q_2^1$	=	-1.78 nT	$q_2^{0,\text{GSM}}$	=	1.43 nT
$s_2^1$	=	0.81 nT			
$q_2^2$	=	-0.58 nT			
$s_2^2$	=	0.07 nT			

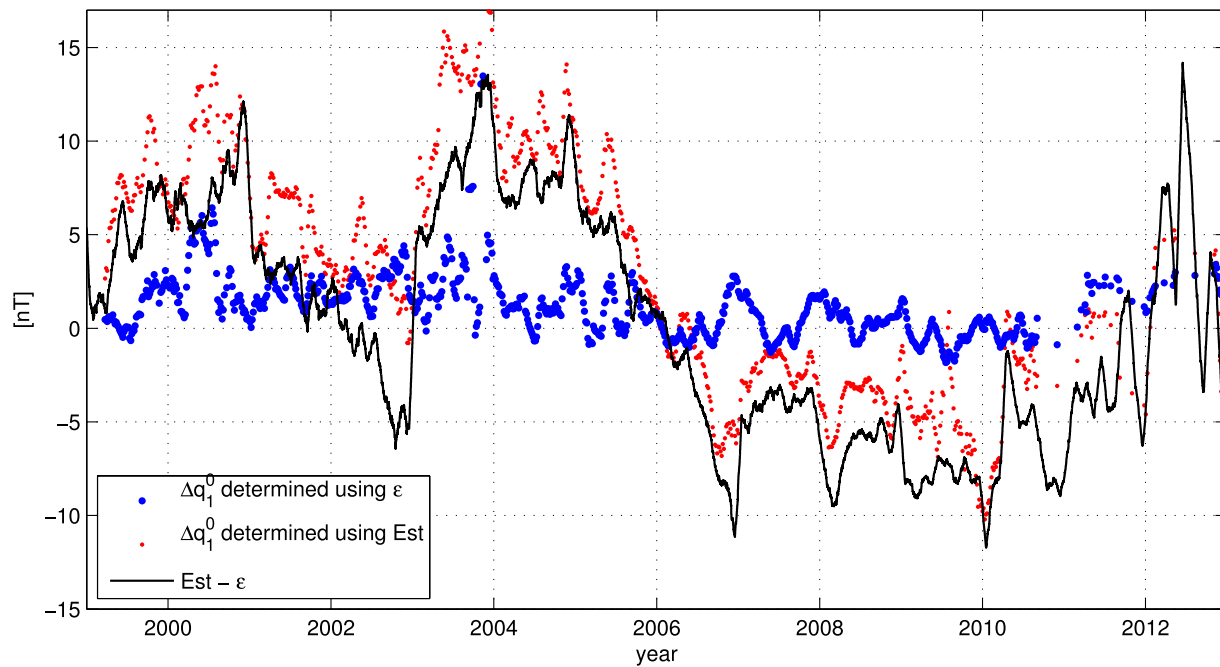
### 4.3 Magnetospheric field contribution

The model coefficients describing the magnetospheric part of CHAOS-4 (cf. eqs 4 and 5) are listed in Table 2.  $\Delta q_1^0 = 1.90$  nT,

the mean value of the external axial dipole terms  $\Delta q_1^0(t)$  in the SM frame averaged over the 900 5-d bins, is much smaller than the corresponding term  $q_1^{0,\text{GSM}} = 11.89$  nT in the GSM frame. If  $RC$  exactly described the magnetic field contribution in the SM frame as seen by the satellites Ørsted, CHAMP and SAC-C one would expect a vanishing ‘baseline correction’  $\Delta q_1^0$  and a regression coefficient  $\hat{q}_1^0 = -1$ ; the estimated values ( $\Delta q_1^0 = 1.90$  nT,  $\hat{q}_1^0 = -0.91$ ) indicate that  $RC$  provides a rather good description of the SM magnetic field contribution at satellite altitude. However, although  $\hat{q}_1^0$  is close to  $-1$  it is nevertheless smaller (in absolute value) than the expected value. This means that the time changes of the SM magnetospheric field as seen by the satellites are only 91 per cent of that predicted by  $RC$  as derived from ground data. Nonetheless, this value is closer to  $-1$  than was the case for previous estimates using  $Dst$  instead of  $RC$  to describe the magnetospheric field contributions.

We also derived a model similar to CHAOS-4I but using  $Dst(t) = Est(t) + Ist(t)$  instead of  $RC(t) = \epsilon(t) + \iota(t)$ . The estimated regression coefficient of that model is  $\hat{q}_1^0 = -0.84$ , which means that the SM field seen by the satellites is only 84 per cent of the one predicted by  $Dst$ . Perhaps even more important is the ‘baseline correction’  $\Delta q_1^0(t)$  of that model, shown by the red dots in Fig. 12. The values scatter considerably more (rms = 6.7 nT) compared to the ones derived using  $RC$  (blue dots, rms = 2.8 nT). This is indicative of the baseline-instabilities of the  $Dst$ -index, as is also obvious when looking at the difference between the external parts of  $Dst$  and  $RC$  (black curve of Fig. 12). All these findings indicate that  $RC$  provides an improved description of the magnetospheric field (in particular during quiet conditions) compared to  $Dst$ .

The non-axial dipole external field coefficients ( $q_1^1, s_1^1$ ) of CHAOS-4 are uncorrelated with  $RC$ , as can be seen from the vanishing regression coefficients  $\hat{q}_1^1, \hat{s}_1^1$ . The mean values  $\Delta q_1^1 = -1.05$  nT,  $\Delta s_1^1 = -0.17$  nT of the corresponding ‘baseline corrections’ are also close to zero, and their rms scatter around the mean is below 1 nT. This indicates that the dipole-part of the magnetospheric field can be well described by the  $RC$  index plus the



**Figure 12.** External field ‘baseline corrections’  $\Delta q_1^0$  in dependence on time for CHAOS-4 (blue dots). Values from a similar model run parametrized by *Dst* instead of *RC* are shown in red. The difference  $\epsilon(t) - Est(t)$  between the external parts of *RC* and *Dst* is shown by the black curve. All values are 30-d running means (see Section 4.3).

field of a degree-one zonal coefficient in the GSM frame of amplitude  $\approx 12$  nT. For comparison Lühr & Maus (2010) reported for the scaling factor  $\hat{q}_1^0 = -0.87$  and the stable GSM field  $q_1^{0,GSM} = 8.31$  nT.

Let us finally discuss the remaining non-dipole external field terms: The largest magnetospheric coefficient (apart from  $q_1^{0,GSM}$ ) is  $q_2^1 = -1.78$  nT. Together with the corresponding coefficient  $s_2^1 = +0.81$  nT it produces a diurnal magnetic field variation of strength 1.9 nT with maximum (negative) amplitude at 22:20 MLT. Although this value has been derived from nightside data only it is in good agreement with the prediction of the Tsyganenko model (e.g. Tsyganenko 1990). Olsen (1996) found a contribution to the degree-2 order-1 spherical harmonic coefficients for geomagnetic quiet conditions ( $Kp = 0$ ) of strength 1.4 nT with maximum (negative) amplitude at 23:48 MLT.

## 5 CONCLUSIONS

Using more than 14 yr of continuous satellite data, augmented with revised monthly means from magnetic observatories, we have derived a model of the static and time-varying part of the Earth’s magnetic field. The model describes rapid core field variations occurring over only a few months as well as the lithospheric field at least up to spherical-harmonic degree  $n = 85$ .

Data from the *Swarm* satellites launched in November 2013 will make it possible to extend the CHAOS model series both regarding better resolution of small-scale lithospheric field features and more rapid time changes of the core field. Processing, calibration and validation of the *Swarm* data during the first months of the mission lifetime requires an accurate reference model of the geomagnetic field which we believe CHAOS-4 can provide. However, one has to be careful when extrapolating CHAOS-4 to times beyond the data span period (1997.5–2013.5). In addition, some model coefficients describing the core field after mid-2010 might be affected by the

Backus effect since only scalar satellite data have been used after 2010 September. This mainly concerns sectorial terms  $g_n^n, h_n^n$ . Inclusion of vector data from *Swarm* will mitigate these problems. We plan to update CHAOS-4 by augmenting with *Swarm* data as soon as these data become available.

Previous versions of CHAOS and the data used for their determination have been widely used for scientific studies, and we hope that CHAOS-4 will also prove useful to the scientific community, both regarding investigations of the dynamics of the core field, the detailed structures of the lithospheric field and in unravelling the various magnetospheric contributions.

Model coefficients of CHAOS-4 and data sets used for their determination are available in different formats at [www.spacecenter.dk/files/magnetic-models/CHAOS-4/](http://www.spacecenter.dk/files/magnetic-models/CHAOS-4/).

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